



# Near-infrared Spectroscopy of Five Ultra-massive Galaxies at $1.7 < z < 2.7$

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## Abstract

We present the results of a pilot near-infrared spectroscopic campaign of five very massive galaxies ( $\log(M_*/M_\odot) > 11.45$ ) in the range of  $1.7 < z < 2.7$ . We measure an absorption feature redshift for one galaxy at  $z_{\text{spec}} = 2.000 \pm 0.006$ . For the remaining galaxies, we combine the photometry with the continuum from the spectra to estimate continuum redshifts and stellar population properties. We define a continuum redshift ( $z_{\text{cont}}$ ) as one in which the redshift is estimated probabilistically from the combination of catalog photometry and the observed spectrum using EAZY. We derive the uncertainties on the stellar population synthesis properties using a Monte Carlo simulation and examine the correlations between the parameters with and without the use of the spectrum in the modeling of the spectral energy distributions. The spectroscopic constraints confirm the extreme stellar masses of the galaxies in our sample. We find that three out of five galaxies are quiescent (star-formation rate of  $\lesssim 1M_\odot \text{ yr}^{-1}$ ) with low levels of dust obscuration ( $A_V < 1$ ), that one galaxy displays both high levels of star formation and dust obscuration ( $\text{SFR} \approx 300M_\odot \text{ yr}^{-1}$ ,  $A_V \approx 1.7$  mag), and that the remaining galaxy has properties that are intermediate between the quiescent and star-forming populations.

*Key words:* galaxies: high-redshift – galaxies: stellar content

## 1. Introduction

The standard  $\Lambda$ CDM paradigm of structure formation depicts a universe dominated by dark matter. In this paradigm, dark matter structures form hierarchically, with low-mass halos forming before their higher mass counterparts. It is therefore expected that galaxies would also form in a hierarchical manner. However, archaeological studies of local galaxies show that today's most massive galaxies formed the bulk of their stars rapidly in the very early universe (see, for example, Thomas et al. 2005): these observations are supported by the existence of a significant population of massive quiescent galaxies up to redshift  $z \sim 3$ . (Cimatti et al. 2002; Kriek et al. 2006; Mancini et al. 2009; Marchesini et al. 2010, 2014; Brammer et al. 2011; Newman et al. 2015; Hill et al. 2016).

Recent models of galaxy formation have been successful in addressing several long-standing tensions between theoretical predictions of massive galaxy growth in the early universe and observational results of such massive galaxies at high- $z$ . These models predict that massive ellipticals will possess old, metal-rich stellar populations and the shortest formation timescales (de Lucia et al. 2006; de Lucia & Blaizot 2007), but there are

still persistent differences between observations and theoretical predictions of the high-mass end of the stellar mass function. These issues raise questions about the evolution of massive galaxies: Marchesini et al. (2014) propose a revised evolutionary path for the formation of today's ultra-massive galaxies (UMGs), in which the progenitor population at  $1.5 < z < 2.5$  is dominated by heavily dust-obscured ( $A_V \sim 2$  mag) massive star-forming galaxies, and only  $\sim 40\%$  of the population is comprised of quiescent galaxies with relatively little dust obscuration.

Though recent photometric campaigns have been highly successful in probing the evolutionary path of UMGs, little is known about the evolution of the most massive galaxies (i.e.,  $\log(M_*/M_\odot) > 11.5$ ) in the early universe. According to measurements of the stellar mass function, the number density of such galaxies appears constant from  $z \sim 4$  to  $z \sim 1.5$  (Marchesini et al. 2009; Muzzin et al. 2013b). In this stellar mass regime, model predictions are, at best, marginally consistent with these observations (Vogelsberger et al. 2014; Davé et al. 2016; but see also Henriques et al. 2015).

Due to the rarity of these galaxies, however, a huge amount of data and manpower are needed to generate a large enough sample to characterize these populations. Very few high- $z$

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galaxies in this mass regime have been confirmed spectroscopically, and population studies rely almost entirely on photometry for redshifts and stellar population characteristics. Recent studies have furthered the spectroscopic effort for estimating the nature of highly massive galaxies at high redshift (Onodera et al. 2010, 2012; van de Sande et al. 2011, 2013; Bezanson et al. 2013; Belli et al. 2014, 2015; Marsan et al. 2015, 2016), but the total sample of spectroscopically confirmed very massive, high- $z$  galaxies known to the community at present is still not large enough to be able to leverage satisfactory statistical power in characterizing the underlying distribution that governs the characteristics of such galaxies.

In an ongoing effort to expand the sample of spectroscopically confirmed very massive galaxies at high- $z$ , we present the results of near-infrared (NIR) spectroscopy of five highly massive galaxies at  $1.7 < z < 2.7$ . In Section 2, we outline the sample selection process. The spectroscopic observations are detailed in Section 3. We derive robust redshifts and stellar population properties using a combination of the NIR spectroscopy and UltraVISTA photometry in Section 4 in order to confirm their status as ultra-massive and to study the characteristics of this rare class of galaxies in the early universe. We present spectroscopic redshifts and modeling results in Section 5. We place the characteristics derived for this sample of high- $z$  ultra-massive galaxies within the context of the population of such galaxies in Section 6. In this work, we assume a Chabrier (2003) IMF and adopt a cosmology of  $\Omega_M = 0.3$ ,  $\Omega_\Lambda = 0.7$ , and  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ .

## 2. The Sample

The targets of this study were chosen from the sample of ultra-massive galaxies identified in the multiwavelength catalog constructed by Muzzin et al. (2013a) across  $1.62 \text{ deg}^2$  of the UltraVISTA/COSMOS field. The UltraVISTA survey itself observes across four broadband filters,  $Y$ ,  $J$ ,  $H$ , and  $K_s$ , as well as one narrowband filter centered on  $H\alpha$  at  $z = 0.8$ , which is not used in the present work (McCracken et al. 2012). The NIR coverage of UltraVISTA probes the rest-frame optical for galaxies at  $1.5 < z < 4.0$ , an important range for the estimation of stellar population properties. The catalog from which the present sample was selected was constructed using the first data release of the UltraVISTA survey; this data release spans about one season of observing time, with a  $5\sigma$  depth of approximately  $K_s < 23.9 \text{ AB}$  in a  $2''$  aperture.

The UltraVISTA catalog spans 30 filters from  $0.15 \mu\text{m}$  to  $24 \mu\text{m}$ , including ultraviolet (UV) imaging from the *GALEX* satellite (Martin et al. 2005) and infrared coverage from *Spitzer* (Sanders et al. 2007; Frayer et al. 2009). Sources are selected from the  $K_s$  band, which reaches 90% completeness at  $K_{s,\text{TOT}} = 23.4 \text{ AB}$ , corresponding to a stellar mass 90% completeness limit of  $\log(M_*/M_\odot) = 10.43$  at  $z = 2$ , as shown in Figure 1. This broad wavelength coverage in the UltraVISTA/COSMOS field allows us to confidently identify high- $z$  massive galaxies candidates for spectroscopic follow-up. Figure 1 shows the stellar mass ( $M_*$ ) versus redshift ( $z$ ) and  $K_s$  magnitude ( $K_{s,\text{tot}}$ ) versus  $M_*$  diagrams, with the five galaxies targeted by our spectroscopic program highlighted as colored filled circles.

We originally selected all galaxies at  $1.7 < z < 2.7$  with  $\log(M_*/M_\odot) > 11.5$ , for a total of 18 candidates spanning the range in  $K_s$ -band total magnitude of  $19.4 < K_{s,\text{tot}} < 23.0$ . From this sample, we selected all galaxies (eight in total)

brighter than  $K_{s,\text{tot}} = 20.6$  to observe spectroscopically. Due to scheduling constraints, two of these galaxies were not observed. The exclusion of these two galaxies does not introduce any additional biases. Therefore, the six galaxies presented in our work constitute a representative sample of the population of brightest and most massive galaxies in the redshift range  $1.5 \lesssim z \lesssim 3.0$ .

Table 1 gives the R.A. and decl. of each of the sources in our sample along with observing and exposure information. Table 2 presents the photometry from selected bands in the UltraVISTA catalog for each of the targeted galaxies.

## 3. Data

Four of the five galaxies in the sample (COS-75355, COS-90676, COS-207144, and COS-189962) were observed with the Gemini Near-infrared Spectrograph (GNIRS; Elias et al. 2006); the remaining two galaxies were observed with the Folded-port InfraRed Echellette (FIRE; Simcoe et al. 2013) mounted on the Magellan 1 telescope. All of the galaxies in our sample were additionally imaged with *Hubble Space Telescope* (*HST*) WFC3 in the F160W band (GO-12990; PI: Muzzin).

### 3.1. GNIRS

Of the four galaxies observed with GNIRS, one (COS-189962) was discarded from the sample due to poor seeing during the observations.

GNIRS was set in its cross-dispersed mode with a 32 line/mm grating and a  $0''.675$  slit during the observing runs, which spanned six dates in March, May, and December of 2013. The observations were made in queue mode, and are taken in 600 s exposures; the details of the observation runs are found in Table 1. The telluric emission lines present in the NIR place a limit on the per-exposure observation time; we therefore opt to take many shorter exposures and subsequently stack the resulting spectra in order to boost the signal-to-noise ratio (S/N). The adopted GNIRS configuration provides a resolution of  $R \approx 800$  and a wavelength range of around  $0.8\text{--}2.4 \mu\text{m}$ . A blind offset star and a B9V star were also observed for acquisition and to derive a telluric correction. Seeing averaged around  $1''.0$  across the observing run, with minimal cloud cover (photometric conditions).

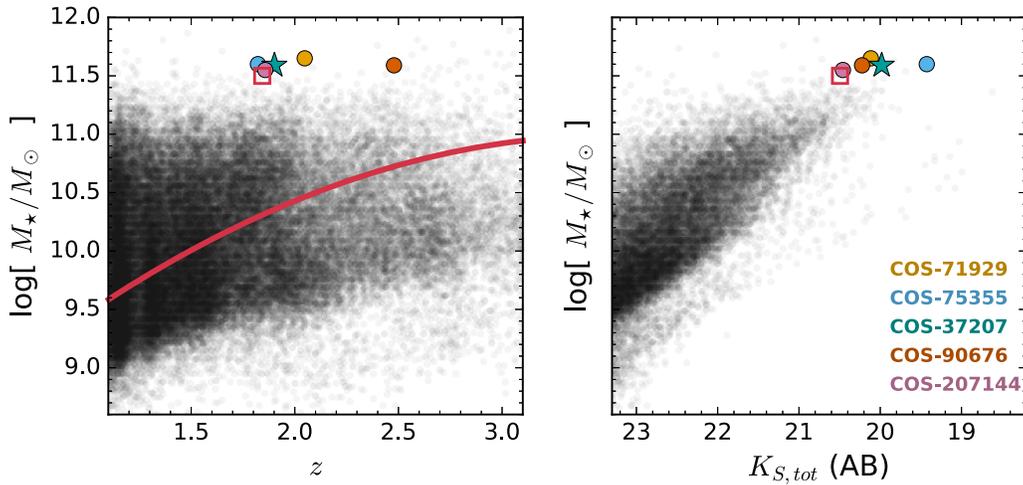
### 3.2. FIRE Observations

FIRE observations were made in its echellette mode with a  $0''.6$  slit and practical coverage from approximately  $1.0 \mu\text{m}$  to  $2.0 \mu\text{m}$  with a spectral resolution of  $R \approx 6000$ .

Observations were taken across two nights in May of 2013. As was the case for the GNIRS observations, several short exposures were taken of each target at 908.8 s per exposure, due to the telluric emission lines present in the observed wavelength range. An A0V star was also observed in order to perform telluric corrections. Seeing averaged between  $0''.7$  and  $1''.0$  throughout the observing run, with minimal cloud cover.

### 3.3. HST Imaging

The galaxies in hand span a significant variety of morphologies, as shown in Figure 2: COS-75355 appears to be consistent with a galaxy undergoing one to two minor mergers and COS-37207 is consistent with being a massive face-on red spiral, while COS-71929 and COS-90676 are relatively isolated with some diffuse emission. For a quantitative discussion of these



**Figure 1.** Left: stellar mass vs. redshift for the present sample (filled star and circles), COS-189962 (unfilled red square, not included in sample due to poor observing conditions), and the UltraVISTA DR1 catalog of Muzzin et al. (2013a) as a whole (grayscale representation). The symbol type represents the classification of the galaxy as star-forming (star) or quiescent (circle) based on the *UVJ* diagram (see Figure 8). The red curve denotes the 90% mass completeness level for the catalog. Right: the same for stellar mass vs.  $K_s$  magnitude for galaxies at  $1.7 < z < 2.7$ . The sample at hand was selected to contain the brightest and most massive galaxies in our targeted redshift range.

**Table 1**  
Pointing and Observation Information for the Initial Sample

id	R.A.	Decl.	Dates	Exp. Time (s) (exposure time × exposure count)	Instrument
COS-75355	10 <sup>h</sup> 02 <sup>m</sup> 28 <sup>s</sup> :49	+02 <sup>d</sup> 02 <sup>m</sup> 13 <sup>s</sup> :70	12/18/13	600 × 10 = 6000	GNIRS <sup>a</sup>
COS-207144	10 <sup>h</sup> 00 <sup>m</sup> 33 <sup>s</sup> :48	+02 <sup>d</sup> 28 <sup>m</sup> 54 <sup>s</sup> :74	03/07/13	600 × 18 = 10800	GNIRS <sup>a</sup>
COS-90676	10 <sup>h</sup> 01 <sup>m</sup> 57 <sup>s</sup> :00	+02 <sup>d</sup> 16 <sup>m</sup> 12 <sup>s</sup> :14	12/17/13	600 × 6 = 3600	GNIRS <sup>a</sup>
COS-189962 <sup>b</sup>	10 <sup>h</sup> 02 <sup>m</sup> 14 <sup>s</sup> :42	+02 <sup>d</sup> 35 <sup>m</sup> 11 <sup>s</sup> :92	05/22/13, 05/23/13	600 × 18 = 10800	GNIRS <sup>a</sup>
COS-71929	10 <sup>h</sup> 01 <sup>m</sup> 40 <sup>s</sup> :60	+01 <sup>d</sup> 58 <sup>m</sup> 57 <sup>s</sup> :47	05/04/13	908.8 × 4 = 3635.2	FIRE <sup>c</sup>
COS-37207	09 <sup>h</sup> 59 <sup>m</sup> 42 <sup>s</sup> :59	+01 <sup>d</sup> 55 <sup>m</sup> 01 <sup>s</sup> :55	05/03/13, 05/04/13	908.8 × 8 = 7270.4	FIRE <sup>c</sup>

**Notes.**

- <sup>a</sup> The Gemini Near-Infrared Spectrograph (GNIRS; Elias et al. 2006).
- <sup>b</sup> This galaxy was omitted from analysis due to poor observing conditions.
- <sup>c</sup> The Folded-port InfraRed Echellette (FIRE; Simcoe et al. 2013).

**Table 2**  
Photometry from the UltraVISTA Catalog for the Six Galaxies in the Initial Sample

filter	UltraVISTA DR1 ID					
	75355	207144	90676	189962	71929	37207
$m_u^*$	25.7	26.4	28.82	24.4	26.7	25.2
$m_{g^+}$	24.8	26.3	26.8	24.3	26.1	24.7
$m_{r^+}$	23.9	26.0	24.7	24.0	25.4	24.1
$m_{i^+}$	23.0	25.1	24.2	23.7	24.4	23.5
$m_{z^+}$	22.1	24.5	23.8	23.2	23.8	22.9
$m_Y$	21.5	23.3	23.1	22.9	22.9	22.5
$m_J$	20.3	22.1	22.1	21.7	21.7	21.3
$m_H$	19.8	21.0	20.6	21.0	20.6	20.5
$m_{K_s}$	19.4	20.5	20.2	20.5	20.1	20.0

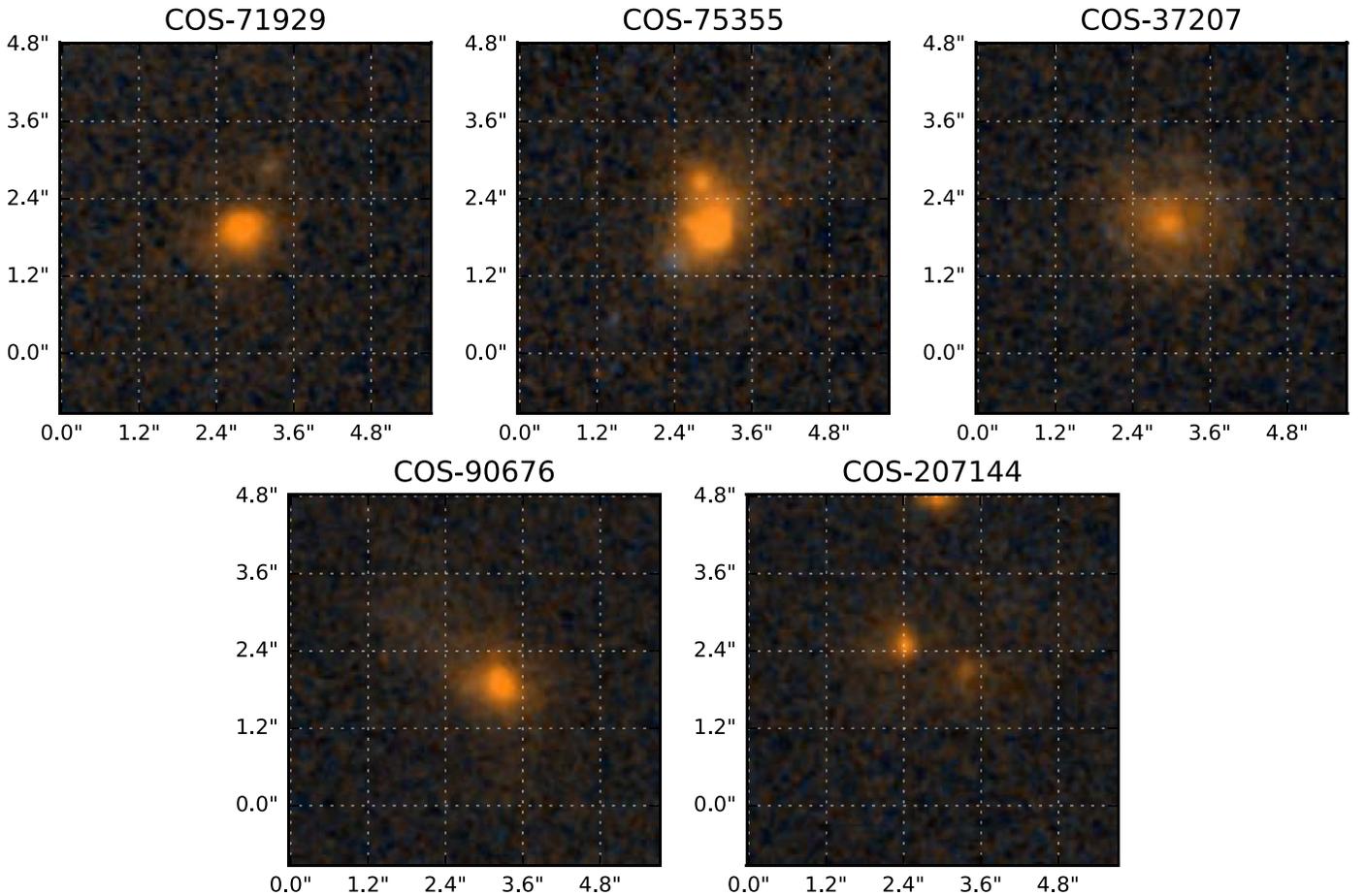
**Note.** Total magnitudes are used. Magnitudes are given in the AB system.

morphologies, see Z. C. Marsan et al. (2017, in preparation). COS-207144, though identified as a single object in the UltraVISTA catalog, is clearly resolved in two components in the *HST* image (see Figure 2). COS-207144 is the only object in

our sample that falls within the 3D-HST survey (Skelton et al. 2014). In 3D-HST, COS-207144 is resolved as two objects, consistent with being an interacting pair at  $z = 2.02^{+0.09}_{-0.14}$  and  $z = 2.07 \pm 0.05$ . There exists a grism redshift for the latter object at  $z_{\text{grism}} = 2.362 \pm 0.015$  (Brammer et al. 2012; Momcheva et al. 2016). However, this grism redshift is based upon a low-S/N, likely spurious, feature identified as [O II] which, if adopted, represents a shift of  $>5\sigma$  from the original photometric redshift, implying that the two resolved objects are not physically interacting. We therefore choose to compare the results of our analysis with those obtained using the original photometric redshift from Skelton et al. (2014).

**3.4. Data Reduction and Extraction of Spectra**

We performed the initial steps of the GNIRS data reduction (cosmic-ray detection, bias subtraction, flat-fielding) using *pyraf*. We then performed a first-pass sky subtraction by constructing a sky frame using up to four dithered exposures taken at sufficiently similar times as the target exposure. We then ran a second pass sky subtraction by using the IRAF task *background* on the target exposure, which fits the sky on each



**Figure 2.** Color cutouts of the galaxies in the present sample constructed from ACS I814 (F814W, blue) and WFC3 H160 (F160W, red) images.

side of the galaxy continuum in order to remove remaining telluric emission lines present in the frame.

To extract the spectra to one dimension, we used the spectrum of the telluric star as a trace for any remaining non-linearity in the spectrum shape. We then performed a telluric correction using the one-dimensional spectrum of the observed telluric star. The orders (GNIRS’s cross-dispersed mode yields six partial spectra corresponding to different wavelength ranges) were then concatenated using a weighted average and flux-calibrated using the UltraVISTA  $H$  and  $K_s$  photometry.

The IDL-based reduction pipeline FIREHOSE (Matejek & Simcoe 2012) was used to reduce the FIRE spectra. We adopted manual tracing for these spectra to account for the faintness of the targets, but left the pipeline otherwise intact.

In order to increase the S/Ns of these observations, we binned each spectrum with bins of variable size. As a result, bin size is not constant across a given spectrum—areas of low S/N are binned more as compared to areas in which the S/N was relatively high in the unbinned spectrum. Areas of low transmission, strong skyline residuals, or artifacts from remaining skyline residuals were masked. The resulting bin size (observed frame) ranges from 51 Å (15th percentile) to 163 Å (85th percentile), with a median bin size of 77 Å. The resulting S/N ranges from 2.51 (15th percentile) to 13.34 (85th percentile), with a median S/N of 6.99.

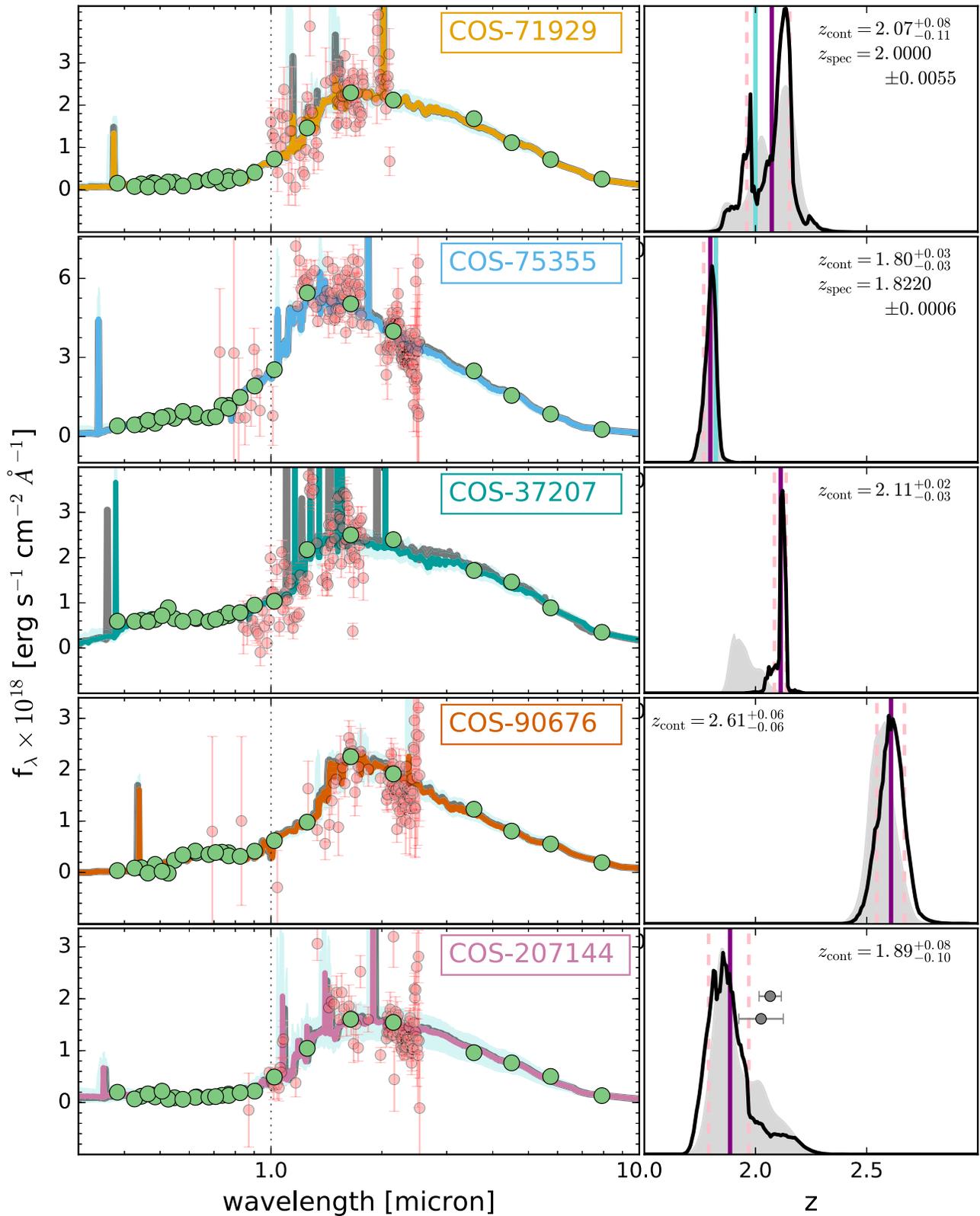
Figure 3 shows the final binned 1D spectra of the five sources (red filled circles), overplotted with the UltraVISTA broadband and medium-band photometry (green filled circles).

#### 4. SED Modeling

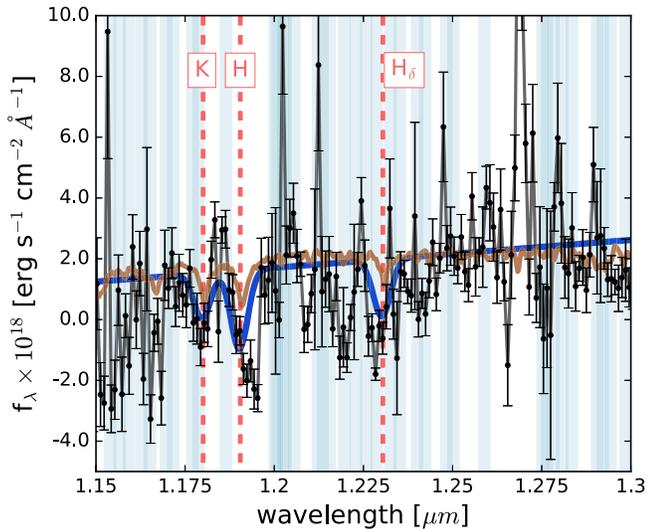
We combined the binned spectra with the UltraVISTA photometry to model the spectral energy distributions (SEDs) of galaxies using the photometric redshift code EAZY (Brammer et al. 2008) and the stellar population synthesis code FAST (Fitting and Assessment of Stellar Templates; Kriek et al. 2009). The addition of the spectra allows us to derive continuum redshifts and stellar population properties from the two codes by taking advantage of the coverage provided by the broadband photometry and the finer observation of the 4000 Å break delivered by the spectroscopy, producing robust redshift estimates in the absence of emission or absorption features.

Following Muzzin et al. (2013b), EAZY fits a linear combination of six templates from the PEGASE models (Fioc & Rocca-Volmerange 1997), a red template from the model library of Maraston (2005), a  $\sim 1$  Gyr old post-starburst template, as well as a slightly dust-reddened Lyman Break template. EAZY also incorporates a template error function to account for larger uncertainties in the models in the rest-frame NIR relative to the rest-frame optical. A detailed description of this process may be found in Brammer et al. (2008). EAZY was run using the v1.0 template error function and a  $K$ -band magnitude prior.

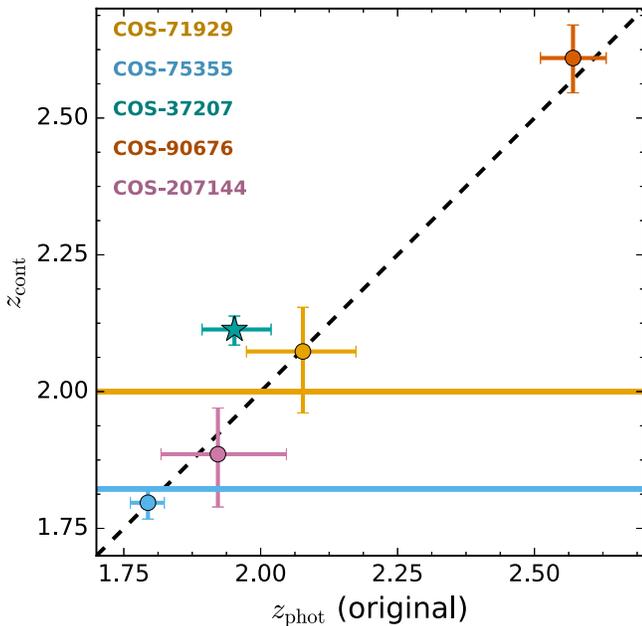
We used FAST to estimate stellar population properties. FAST produces best-fit parameters for stellar mass, star-formation rate, stellar age, characteristic timescale ( $\tau$ ) of the assumed star-formation history (i.e., delayed exponentially



**Figure 3.** Left: binned spectra for each galaxy (red filled circles) with the best-fit EAZY model (color corresponds to the galaxy’s label in each panel). Fluxes are scaled by a factor of  $10^{18}$ . The light blue shaded region encases the upper and lower envelopes of the EAZY models that correspond to redshifts within the 68% confidence bound of the redshift posterior distribution. The green filled circles show the UltraVISTA catalog photometry. Right: EAZY redshift distributions. The black curve shows the distribution estimated using both the spectrum and photometry. The gray shaded region shows the same distribution from the photometry alone.  $z_{\text{cont}}$  values are shown as purple lines, with the 68% confidence intervals marked with dashed pink lines. The  $z_{\text{spec}}$  of COS-75355 from Onodera et al. (2012) and of COS-71929 from this work are shown as cyan vertical lines. The gray circles in the lower right panel denote the photometric redshifts derived for the two resolved objects that make up COS-207144 in Skelton et al. (2014).



**Figure 4.** Ca II H & K absorption features in COS-71929. The brown curve shows the best-fit EAZY SED to the galaxy at  $z = 2.000$ . The red dashed lines show the placement of the Ca II H & K and  $H_\delta$  features, as labeled, at  $z = 2.000$ . The blue curve shows the fit to a triple Gaussian with the standard deviations and redshifts constrained to be equal for all features. The light blue shaded regions show areas in which OH skylines were removed during data reduction.



**Figure 5.** Continuum EAZY redshifts ( $z_{\text{cont}}$ ) vs. photometry-only EAZY redshifts ( $z_{\text{phot}}$ ). Spectroscopic redshifts are shown for COS-75355 and COS-71929 in blue and yellow horizontal lines, respectively. Quiescent galaxies are marked as circles; the star-forming galaxy COS-37207 is shown by a star. The black dashed line shows the 1:1 relation.

declining star-formation history), extinction (using the extinction curve detailed in Kriek & Conroy 2013), and metallicity. In this analysis, we set an allowed stellar age range of  $\log(\text{age} [\text{yr}]) = 8 - 10.1$ , an allowed range of  $\tau$  of  $\log(\tau [\text{yr}]) = 7 - 10$ , and an allowed extinction range of  $A_V = 0-5$  mag. We used the high-resolution version of the flexible stellar population synthesis (FSPS) models from Conroy et al. (2009) with a Chabrier (2003) IMF.

In order to estimate errors on the stellar population properties derived from FAST, we constructed a Monte Carlo (MC) simulation external to FAST and EAZY. The input photometry and spectra were perturbed according to a normal distribution described by  $\sim N(f_i, \sigma_i)$ , where  $f_i$  is the mean flux of the  $i$ th measurement, and  $\sigma_i$  is the uncertainty of the  $i$ th measurement in the spectrum and photometry. At this juncture, we implicitly assumed that each wavelength bin is uncorrelated with its neighbors; we do expect there to be some degree of correlation between adjacent flux bins in the spectra, but the impact of a bin-to-bin correlation will be an overestimation of the MC-derived errors. We therefore chose to assume that each bin is fully independent of the state of nearby wavelength bins.

For those galaxies that have spectroscopic redshifts, the FAST redshift was held fixed for all runs. For the galaxies that do not have spectroscopic redshifts, we run EAZY using the perturbed spectrum and photometry, and run FAST with the perturbed data, holding the redshift fixed to the best-fit redshift of the corresponding EAZY output.

This method of estimation allows for the independent derivation of errors on the stellar population properties estimated by FAST and EAZY. For all galaxies, 500 simulations were run for both FAST and EAZY. A full description of the parameter correlations that result from these simulations is given in Appendix.

## 5. Results

### 5.1. Spectroscopic Redshifts

There are spectroscopic redshifts for two galaxies in the sample. First, COS-75355 has a literature  $z_{\text{spec}} = 1.822$  from Onodera et al. (2012). Second, the spectrum of COS-71929 shows absorption features consistent with Ca II H & K, allowing us to measure a spectroscopic redshift at  $z_{\text{spec}} = 2.000 \pm 0.006$ .

Onodera et al. (2012) finds a spectroscopic redshift for COS-75355 using the Ca II H & K absorption features, in conjunction with  $H_\beta$  in absorption. The spectrum in our sample does not have sufficient S/N to detect most of the features used to identify the spectroscopic redshift—there is a possible detection of Na in absorption at  $z = 1.822$ , but the detection is not secure.

The features used to determine a spectroscopic redshift for COS-71929 are shown in Figure 4, as labeled. The best-fit EAZY model SED is overlaid in brown. From this spectrum, we constrained our spectroscopic redshift for COS-71929 to be  $z_{\text{spec}} = 2.0000 \pm 0.0055$ , which is within  $1\sigma$  of the continuum redshift, as detailed in the following section.

We simultaneously fit the Ca II H & K and  $H_\delta$  features using a triple Gaussian with the observed velocity dispersion and redshift constrained to be equal for all features. The underlying continuum was approximated as linear in this regime. From this treatment, we derived a formal error for the spectroscopic redshift of  $\pm 0.0055$ . We note that the spectroscopic redshift for COS-71929 is located at the secondary peak in the galaxy's redshift probability density function, with the bimodality of the probability density function likely to be caused by a degeneracy between the locations of CaII H and  $H_\delta$ .

We additionally attempted to approximate a constraint on the stellar velocity dispersion of COS-71929 using the fit detailed above. The instrumental spectral resolution is given to be  $50 \text{ km s}^{-1}$ ; we subtracted this value in quadrature from the observed velocity dispersion to make an estimate of the stellar

**Table 3**  
FAST Best-fit Parameters

id	207144	37207	71929	75355	90676
$A_V$ (mag)	$0.00^{+1.10}_{-0.00}$	$1.70^{+0.10}_{-0.40}$	$1.30^{+0.10}_{-0.40}$	$0.30^{+0.10}_{-0.10}$	$0.50^{+0.30}_{-0.50}$
log (Age [yr])	$9.50^{+0.00}_{-0.60}$	$9.00^{+0.40}_{-0.00}$	$9.50^{+0.00}_{-0.10}$	$9.00^{+0.10}_{-0.10}$	$9.10^{+0.20}_{-0.40}$
log ( $M_*/M_\odot$ )	$11.46^{+0.00}_{-0.26}$	$11.67^{+0.12}_{-0.01}$	$11.87^{+0.00}_{-0.12}$	$11.53^{+0.05}_{-0.02}$	$11.64^{+0.06}_{-0.16}$
log (SFR [ $M_\odot$ yr $^{-1}$ ])	$-0.18^{+0.28}_{-1.82}$	$2.53^{+0.12}_{-0.46}$	$0.94^{+0.02}_{-0.67}$	$-5.41^{+5.80}_{-5.80}$	$-0.54^{+1.12}_{-1.46}$
log (sSFR [yr $^{-1}$ ])	$-11.64^{+0.37}_{-87.36}$	$-9.14^{+0.10}_{-0.55}$	$-10.93^{+0.09}_{-0.60}$	$-16.93^{+5.75}_{-0.00}$	$-12.18^{+1.10}_{-86.82}$
log ( $\tau$ [yr])	$8.50^{+0.13}_{-1.50}$	$8.50^{+0.50}_{-0.00}$	$8.60^{+0.00}_{-0.10}$	$7.60^{+0.40}_{-0.00}$	$8.00^{+0.00}_{-1.00}$
Z	$0.03^{+0.00}_{-0.02}$	$0.01^{+0.00}_{-0.01}$	$0.00^{+0.00}_{-0.00}$	$0.01^{+0.02}_{-0.00}$	$0.00^{+0.03}_{-0.00}$
$z_{\text{peak}}$ or $z_{\text{spec}}$	$1.89^{+0.08}_{-0.06}$	$2.11^{+0.01}_{-0.04}$	$2.0000^{+0.0055}_{-0.0055}$	$1.8220^{+0.0006}_{-0.0006}$	$2.61^{+0.05}_{-0.08}$
$q_{\text{sf}}$	$0.993 \pm 0.005$	$0.276^{+0.529}_{-0.276}$	$0.963 \pm 0.028$	$1.0^{+0.0}_{-0.008}$	$0.999^{+0.001}_{-0.008}$

velocity dispersion. We estimated the uncertainty in the fit due to measurement error by running a Monte Carlo simulation to convergence in which the binned flux values were perturbed according to a normal distribution with a standard deviation equal to the measurement uncertainty. This procedure gives a velocity dispersion of  $549^{+106}_{-262}$  km s $^{-1}$ , where the quoted errors are the two-sided 95% confidence bounds. However, we stress that our measurement of the velocity dispersion is very crude. The low S/N of the spectrum does not allow for a proper analysis of the velocity dispersion, and our measurement is subject to extremely high uncertainty, with a one-sided 99% lower confidence bound of 217 km s $^{-1}$ .

### 5.2. Continuum Redshifts

For all of the galaxies in the sample (including those that have spectroscopic redshifts), we estimated a continuum redshift with EAZY ( $z_{\text{cont}}$ ), in which both the photometry and spectroscopy are used in the fit. Figure 5 shows the continuum redshifts,  $z_{\text{cont}}$ , for the galaxies in our sample against their original photometric redshifts ( $z_{\text{phot}}$ , derived from the Ultra-VISTA DR1 photometry only). The quoted errors in Figure 5 are the 68% confidence intervals from the EAZY redshift probability density function.

The continuum redshifts of both COS-75355 and COS-71929 are in good agreement ( $<1\sigma$  difference) with their spectroscopic redshifts. The spectroscopic redshifts of both galaxies are shown as horizontal lines in Figure 5. Although the error bars shown in the figure are taken from the EAZY redshift probability density function (PDF), the distribution of our EAZY Monte Carlo simulation is in good agreement with the output EAZY PDF. The main peak of the EAZY redshift probability density function is somewhat higher than the spectroscopic redshift for COS-71929 (see Figure 3), but there is a secondary peak in the EAZY redshift PDF that is consistent with our spectroscopic redshift. Furthermore, the spectroscopic redshift is within  $1\sigma$  of the continuum redshift for the galaxy.

The formal errors on the continuum redshift are lower than the original photometric redshift in all cases when the spectrum is added to the fit, although COS-37207 is the only galaxy whose continuum redshift is significantly different from the original photometric redshift. Although the original photometric redshift of COS-37207 lies outside of the 95% confidence interval of the continuum redshift, the absolute difference between the continuum and photometric redshifts is relatively small ( $\sim 0.16$ ; i.e.,  $\Delta z/(1+z) \sim 0.05$ ).

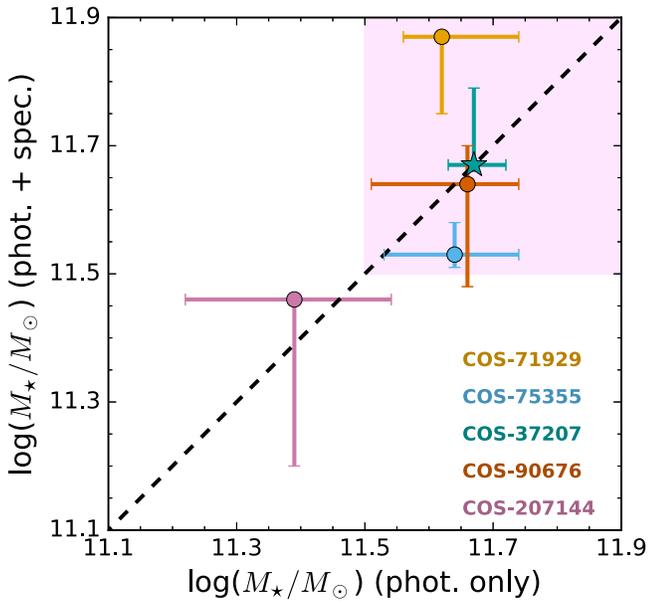
As the photometric redshifts for both galaxies that make up COS-207144 are available, it is informative to compare our results against the photometric redshifts available from Skelton et al. (2014). Given that we expect the majority of the light in the observed spectrum to originate from the brighter galaxy (which, in the NIR, is the galaxy at  $z = 2.02$ ), it is unsurprising that our continuum redshift is in better agreement with the brighter ( $<1\sigma$  difference) than with the fainter ( $<2\sigma$  difference) of the two objects.

### 5.3. Stellar Population Properties

We used the procedure outlined in Section 4 to estimate the stellar population properties of the sample. Table 3 shows the FAST output for the five galaxies, along with their  $1\sigma$  errors. The best-fit values that we quote are the original FAST best-fit values, and not average values of the Monte Carlo distributions. Thus, in some cases, the best-fit value lies on the edge of the  $1\sigma$  interval. Redefining our best-fit values by using an average of the Monte Carlo distributions does not qualitatively change our results.

We confirm that all of the galaxies in our sample are indeed ultra-massive. All galaxies in the sample have a stellar mass exceeding  $\log(M_*/M_\odot) \gtrsim 11.5$ . In the case of COS-207144 (the blended object), stellar mass estimates obtained here may be compared against those obtained for the objects resolved in 3D-HST by Skelton et al. (2014). The best-fit stellar mass of the brighter galaxy of the pair is in good agreement (within the 68% confidence interval) with that of COS-207144. The sum of the best-fit stellar masses of the resolved objects is significantly larger ( $\log(M_*/M_\odot) = 11.62$ ) than our best-fit mass for COS-207144. This is likely a result of the brighter galaxy dominating the light of the spectrum—the best-fit stellar mass of COS-207144 ( $\log(M_*/M_\odot) = 11.46^{+0.00}_{-0.26}$ ) is in strong agreement with the best-fit stellar mass of this source as derived by Skelton et al. (2014) of  $\log(M_*/M_\odot) = 11.40$ .

Given that the original selection criteria from the Ultra-VISTA DR1 catalog was based partially on stellar mass, we plot the stellar mass FAST estimated from photometry alone against the same quantity from the photometry in conjunction with the spectroscopy in Figure 6. None of the updated best-fit stellar masses differ significantly from the stellar mass derived from the photometry alone. In the case of COS-37207, the galaxy for which the best-fit redshift changed the most significantly when the spectrum was added to the fit (see Figure 3), the formal error on the stellar mass increased when



**Figure 6.** Continuum FAST stellar mass estimates vs. photometry-only FAST stellar mass estimates. The black dashed line shows the 1:1 relation. The pink shaded region marks stellar masses where  $\log(M_*/M_\odot) > 11.5$ . As in Figure 5, quiescent galaxies are shown with circles and the star-forming galaxy with a star.

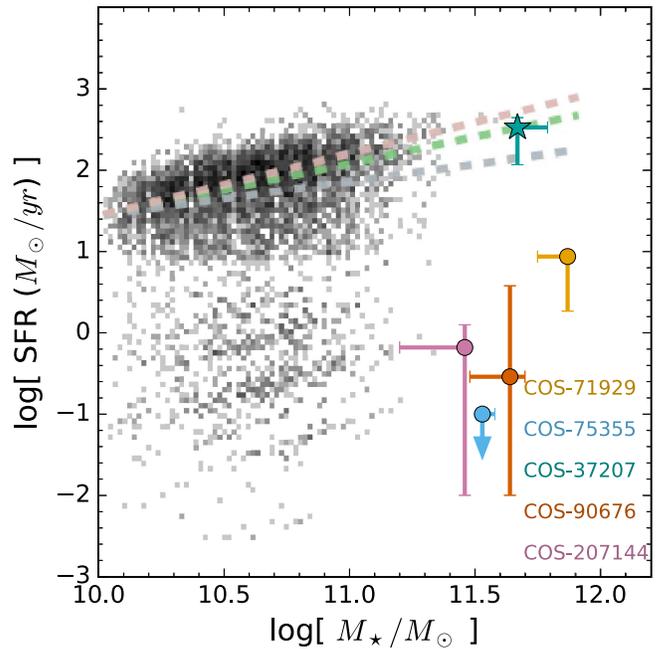
the spectrum is added, with relatively little change to the best-fit stellar mass.

As a probe of the evolutionary stage of the galaxies in our sample, we computed the quenching factor  $q_{\text{sf}}$  (Kriek et al. 2009), which is  $q_{\text{sf}} = 1 - b$ , where  $b$  is the birthrate parameter (Kennicutt 1983; Scalo 1986). The quenching factor is therefore

$$q_{\text{sf}} = 1 - \frac{\text{SFR}_{\text{current}}}{\langle \text{SFR}_{\text{past}} \rangle} = 1 - \frac{\text{SFR}_{\text{current}}}{M_*/\text{age}}, \quad (1)$$

such that a fully quenched galaxy will have a quenching factor of 1. With the exception of COS-37207, which has a quenching factor of  $0.276^{+0.529}_{-0.276}$ , all of the galaxies in our sample have quenching factors consistent with  $q_{\text{sf}} > 0.90$ .

COS-37207 is the only galaxy in the sample to show high levels of star formation, with a best-fit SED star-formation rate (SFR) of  $339^{+108}_{-221} M_\odot \text{ yr}^{-1}$ . The error bars quoted here correspond to the 68% confidence interval. The 99% confidence interval bounds are  $[31, 788] M_\odot \text{ yr}^{-1}$ , the lower bound of which is greater than the best-fit SFR for any of the other galaxies in the sample. We therefore conclude that this galaxy is securely identified as star-forming. In light of the fact that this galaxy has a high estimated SFR, but is also characterized by a large dust extinction of  $A_V = 1.7^{+0.1}_{-0.4}$ , it is of interest to note that the spectrum of COS-37207 does not show any secure emission lines. Using the best-fit values of SFR and  $A_V$ , we estimated the expected observed flux for [O II],  $H_\beta$ , and [O III] using the relations from Calzetti et al. (2000) and Price et al. (2014). From these estimates, we conclude that neither  $H_\beta$  nor [O II] would be observable given the noise properties of the spectrum. [O III] may have been visible, but coincides with a strong skyline at the best-fit redshift of  $z = 2.11$ . It is therefore expected that strong emission lines are absent from the spectrum of COS-37207.

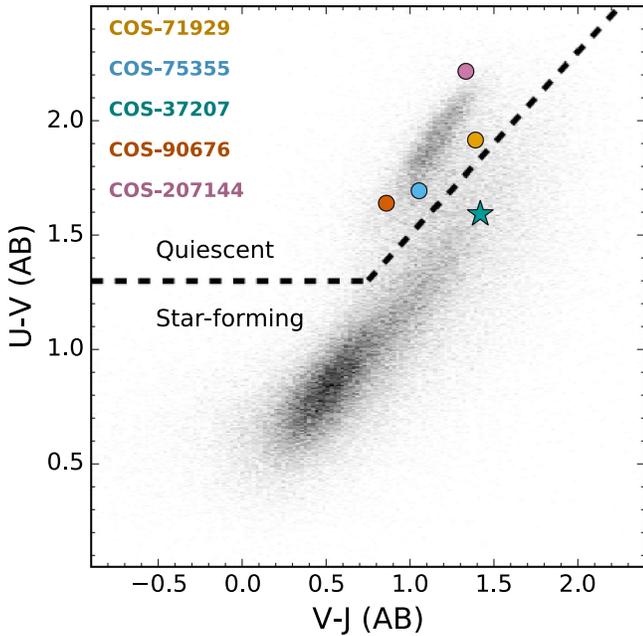


**Figure 7.** Stellar mass vs. star-formation rate for the sample (colored points) and the mass-complete UltraVISTA catalog in the range of  $1.7 < z < 2.7$ . The average star-forming main sequence from Whitaker et al. (2014) for redshift range in question is shown by the green dashed line. The parameterization of the star-forming main sequence by Tomczak et al. (2016) is shown for  $z = 2.15$  by the blue dashed line. The parameterization of the star-forming main sequence as derived by Speagle et al. (2014) at  $t = 3.0$  Gyr is shown by the red dashed line. As expected, the four quiescent galaxies in the sample (filled circles) lie below the star-forming main sequence, whereas COS-37207 (the dusty star-forming galaxy, filled star) lies along the star-forming main sequence.

Although COS-71929 also shows evidence of some star formation ( $\text{SFR} = 8.7^{+0.4}_{-6.9} M_\odot \text{ yr}^{-1}$ ), a measurement of the quenching factor of this galaxy indicates that the SFR was significantly higher in the past ( $q_{\text{sf}} \approx 0.95$ ).

Additionally, the estimated dust extinction for COS-37207 is the highest in the sample, with  $A_V = 1.70^{+0.10}_{-0.40}$  mag. COS-71929 displays the second-highest level of dust extinction at  $A_V = 1.30^{+0.10}_{-0.40}$  mag. We therefore characterize COS-37207 as a dusty star-forming galaxy, COS-71929 as a galaxy with mostly suppressed star formation intermediate between star-forming and quiescent galaxies, and the rest of the galaxies in our sample as quiescent.

We plot our stellar mass and SFR estimates against the mass-complete sample of galaxies at  $1.7 < z < 2.7$  from the UltraVISTA catalog in Figure 7, along with the star-forming main sequences measured by Speagle et al. (2014), Whitaker et al. (2014), and Tomczak et al. (2016). For the UltraVISTA catalog, we preferentially use UV+IR SFR measurements. In the absence of such a measurement (or a non-detection), we use the SED SFR. The FAST-derived stellar masses and SFRs for our sample place the three quiescent galaxies (COS-75355, COS-207144, and COS-90676) on the high-mass end of the quiescent population. COS-37207 lies along the high-mass end of the star-forming main sequence, making it the most massive star-forming galaxy in the targeted redshift range. Finally, COS-71929, with an  $\text{SFR} \approx 10 M_\odot \text{ yr}^{-1}$ , is seen to lie well below the main sequence of star-forming galaxies, indicating that the star-formation activity of COS-71929 is mostly quenched. We note that this picture qualitatively holds if the



**Figure 8.** *UVJ* diagram for the full UltraVISTA DR1 (grayscale) and the present sample (colored points, as labeled). The black dashed lines separate quiescent (top left) from star-forming (right) galaxies. Given the position of COS-37207 in the *UVJ* diagram, this galaxy is characterized as dusty star-forming, in agreement with the results of the SED modeling. The boundaries shown follow those presented in Whitaker et al. (2015) and Martis et al. (2016).

SED-modeling SFRs are replaced with the SFR from the combination of UV (i.e., unobscured) and infrared (dust reprocessed) estimates of the SFRs (see Section 5.5).

All of the galaxies in our sample have nominal best-fit stellar ages greater than or equal to 1 Gyr. Here, stellar ages are given as the age since the onset of star formation. COS-207144 and COS-71929 have best-fit ages comparable to the age of the universe at their redshift, whereas COS-75355 and COS-90676 have best-fit stellar ages a factor of 2–3 younger than the age of the universe at their redshift. The star-forming galaxy COS-37207 is 0.5 dex younger than the age of the universe at  $z = 2.1$ . We note however that the stellar age is typically one of the most uncertain stellar population properties (Muzzin et al. 2009).

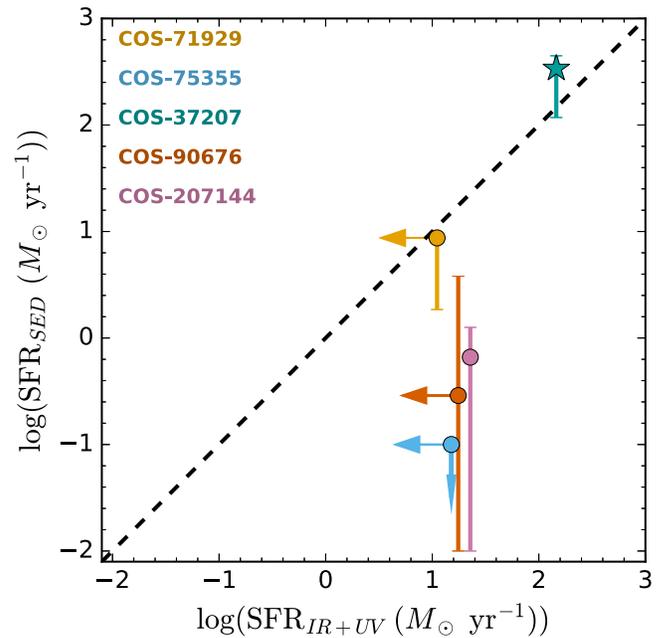
Given the analysis presented in the Appendix and summarized in Section 5.4, the metallicity is not constrained due to degeneracies in the fitted stellar population synthesis parameters.

#### 5.4. SPS Parameter Degeneracies

Because much of this analysis is based upon conclusions reached through the use of SPS modeling, it is imperative to examine the degeneracies between the derived parameters of our SPS models.

In the Appendix, we plot the correlations between EAZY estimates of redshift, FAST estimates of stellar mass, SFR, dust attenuation, and population age for the Monte Carlo simulations of both the photometry only and the photometry + spectrum. We use a Gaussian kernel density estimate in order to visualize the distribution of the output parameters.

As expected, the FAST estimates increase in precision the most when a spectroscopic redshift is present. Although COS-37207 does not have a spectroscopic redshift, its best-fit redshift shifted significantly with the addition of the spectrum (see Figure 3); the

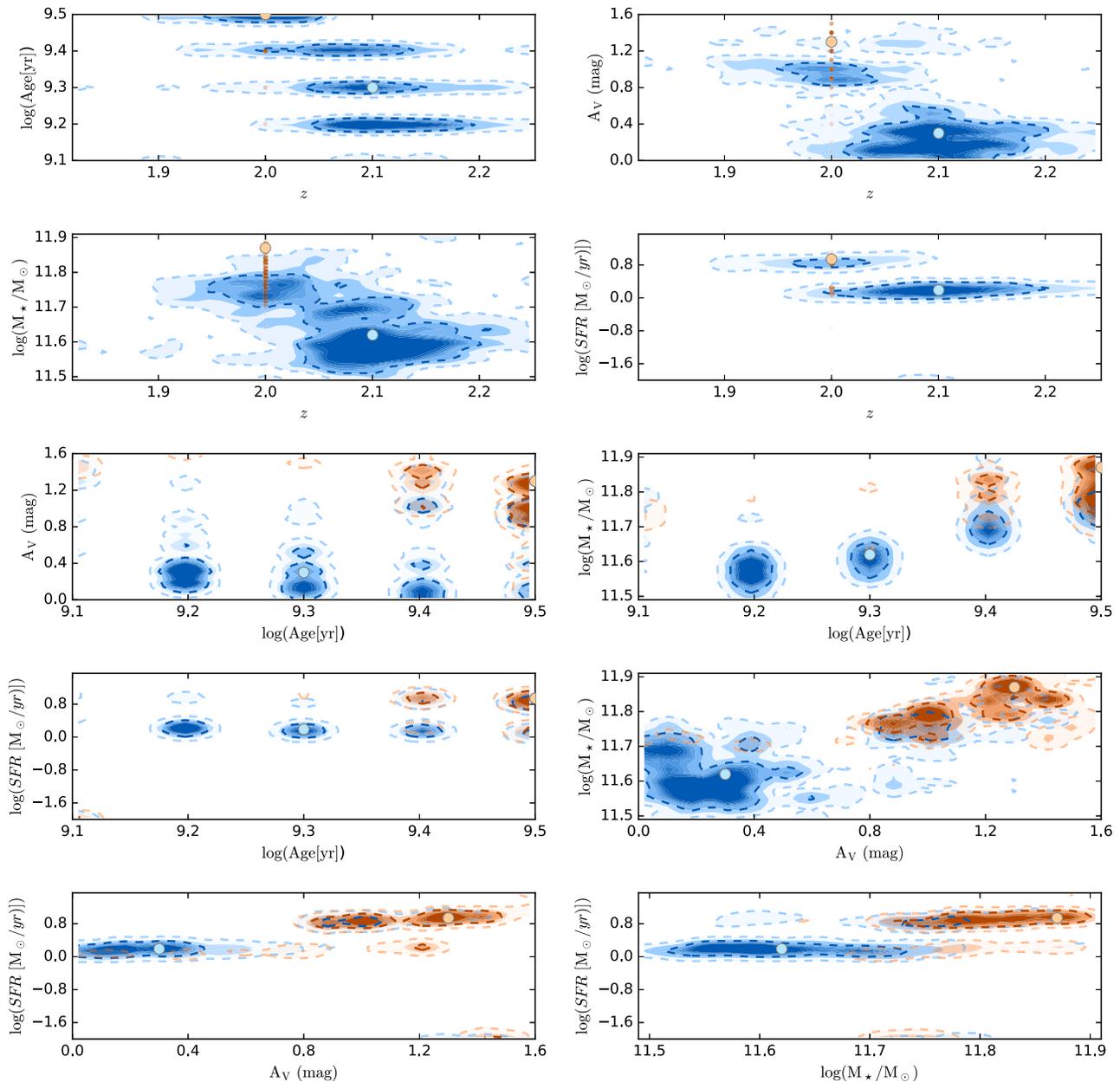


**Figure 9.** Comparison between the SFR estimated from SED modeling and the SFR from the combination of UV and IR measurements. Upper limits are given for the UV+IR SFR when there is a non-detection in the MIPS 24  $\mu\text{m}$  catalog photometry. With the exception of COS-207144, the galaxies in the sample at hand have SFR<sub>SED</sub> estimates that are consistent with the measurement of SFR<sub>UV+IR</sub> from Muzzin et al. (2013a). The black dashed line shows the 1:1 relation.

error on the FAST parameter estimates decreased slightly with the addition of the spectrum, but the best-fit FAST parameter estimates remain relatively unchanged.

As the only vigorously star-forming galaxy in the sample, it is of interest to note that there exists a strong correlation between dust attenuation and SFR in the Monte Carlo distribution of COS-37207. This correlation is not found in any of the other galaxies’ best-fit parameter distributions. This finding does not change the classification of COS-37207 as dusty and star-forming, i.e., even at the lowest dust attenuation and SFR produced by the simulation, COS-37207 would be classified as dusty and star-forming. However, it reaffirms that low-resolution spectra are not sufficient to break the degeneracy between SFR and dust attenuation (Förster Schreiber et al. 2004; Muzzin et al. 2009).

As expected from other similar studies (e.g., Worthey 1994; Muzzin et al. 2009; López Fernández et al. 2016), we find several other correlations between derived stellar population properties. Stellar population age, dust attenuation, and metallicity all redden the overall galaxy SED, and in the absence of prior knowledge that can break degeneracies (Conroy & van Dokkum 2012), it is expected that there will be anticorrelations between these parameters as derived from SED modeling. For all of the galaxies in our sample, we see the expected negative correlation between stellar population age and dust attenuation to some degree. The age–extinction correlation is especially prominent in COS-207144 and COS-75355. This is likely because the addition of the spectrum (and spectroscopic redshift, in the case of COS-75355) eliminated several competing parameter sets that produced discrete populations within the Monte Carlo distribution, as can be seen in the third panel of the left column of Figure 11.



**Figure 10.** Distributions of Monte Carlo best-fit FAST estimates for COS-71929 of the redshift, stellar population age, dust attenuation, stellar mass, and star-formation rate. Monte Carlo distributions are shown for the case when FAST is run with UltraVISTA catalog photometry only (blue) and with the spectrum in conjunction with the photometry (orange). Dotted contours show the 68% confidence (dashed inner, darker curve) and 95% confidence (dashed outer, lighter curve) contours. The blue and orange points represent the best-fit FAST estimate from the original (unperturbed) spectrum for the photometry-only and photometry + spectrum fit, respectively.

### 5.5. $M_*$ and SFR: Comparison with Independent Estimates

In order to assess the accuracy of our previously derived stellar population properties, we compared our results to quiescent/star-forming classifications using the rest-frame  $U-V$  versus  $V-J$  color-color diagram (hereafter  $UVJ$  diagram; following, e.g., Whitaker et al. 2011, 2015; Muzzin et al. 2013b; Marchesini et al. 2014; Martis et al. 2016), and SFRs derived from the UV and infrared (IR), as provided by Muzzin et al. (2013a).

Figure 8 shows the targeted sample in the  $UVJ$  diagram using their original photometric redshifts from Muzzin et al. (2013a). Also plotted is the overall galaxy population in a stellar-mass-complete sample at the same redshift range as that obtained using the UltraVISTA DR1 catalog of Muzzin et al. (2013a).

Figure 8 shows that the location of the targeted galaxies in the  $UVJ$  diagram is consistent with the characterization of their stellar population properties as derived from FAST. In particular, we note that COS-37207 is in the region of the  $UVJ$  diagram typically populated by dusty star-forming galaxies (e.g., Martis et al. 2016), while COS-71929 is in an intermediate zone between the dusty star-forming and quiescent galaxies.

We checked the SED star-formation rates for our sample against star-formation rates derived using  $L_{\text{IR}} + L_{\text{UV}}$  measurements based on the UltraVISTA catalog. We determine the rest-frame UV flux following Muzzin et al. (2013a), estimating  $L_{2800}$  by integrating the best-fit template generated by EAZY (Brammer et al. 2008) from 2600–2950 Å.

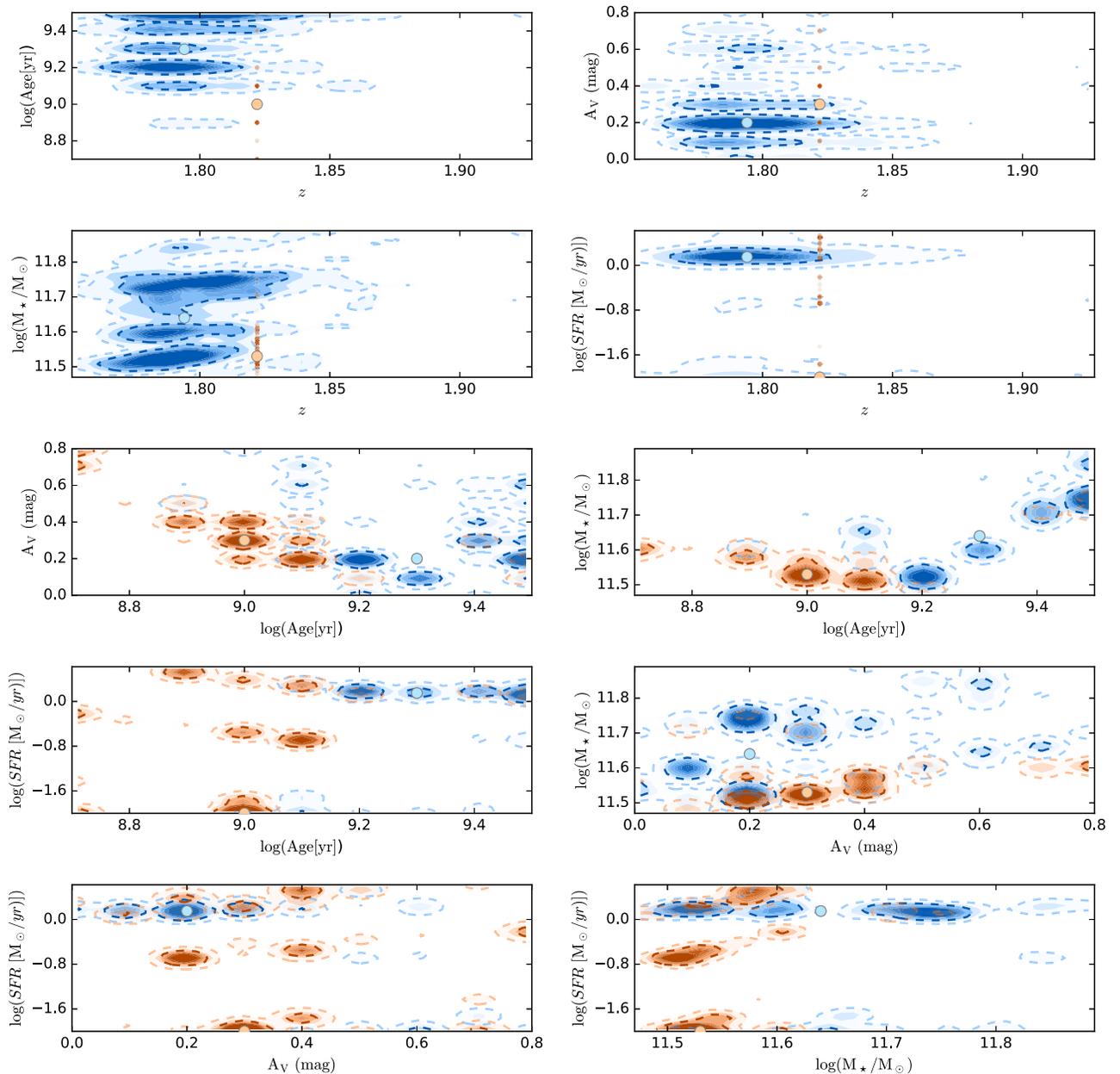


Figure 11. Same as Figure 10 for COS-75355.

Infrared luminosity ( $L_{\text{IR}}$ ) was estimated using the  $24\ \mu\text{m}$  emission alone; we caution that this approach can produce individual measurements of  $L_{\text{IR}}$ , which are uncertain up to several factors. A full explanation of the method to determine  $L_{\text{IR}}$  for the sample at hand can be found in Wuyts et al. (2008); briefly,  $L_{\text{IR}}$  is computed for the Dale & Helou (2002) infrared SEDs of star-forming galaxies at several heating levels of the interstellar environment. The best estimate of  $L_{\text{IR}}$  is then taken to be the log-average of the set of  $L_{\text{IR}}$  estimates for the template across the range of heating levels sampled.

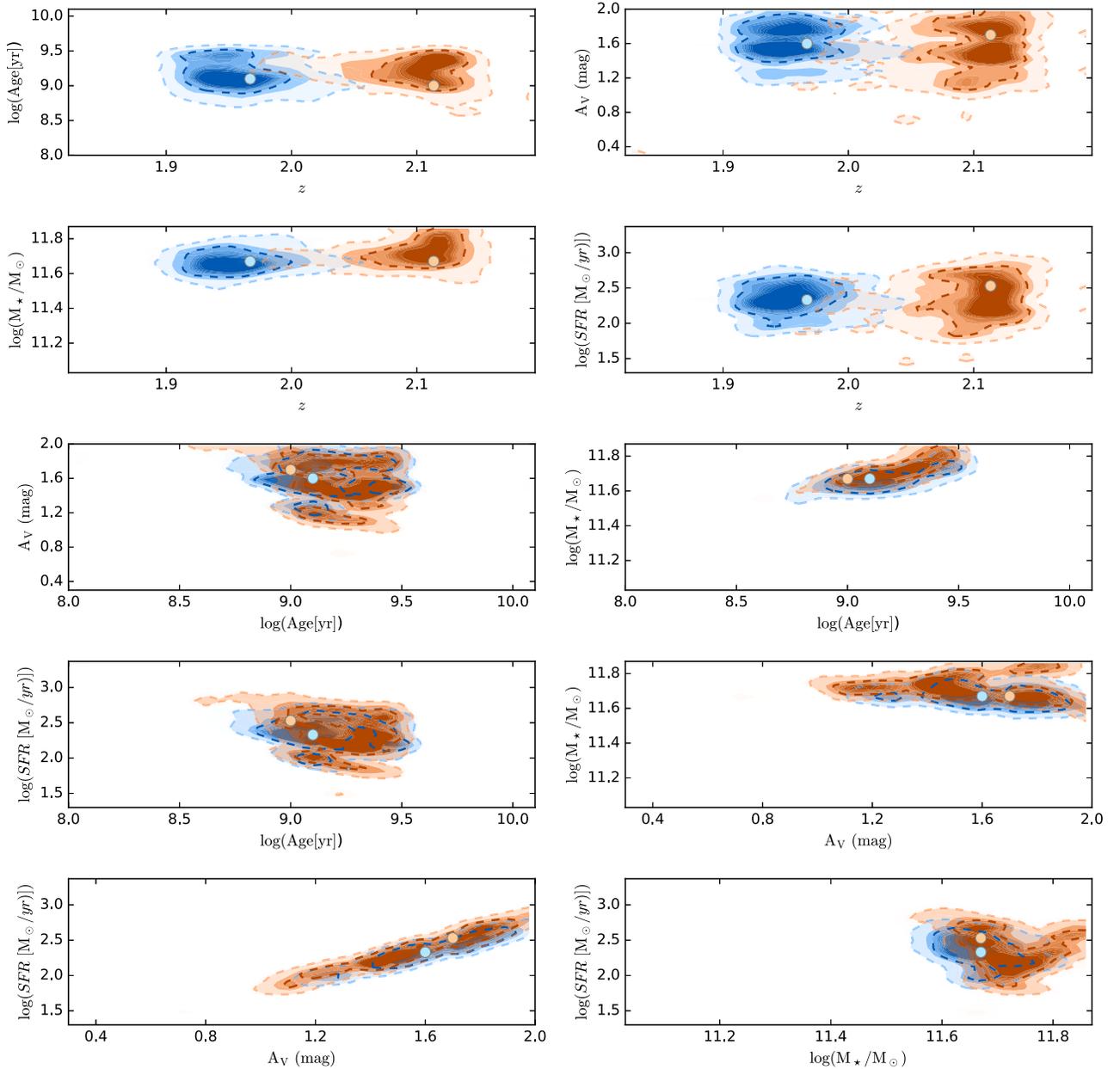
We calculate IR+UV SFRs following Kennicutt (1998) using the calibration of Whitaker et al. (2014), i.e.,  $\text{SFR}_{\text{IR+UV}} = 1.09 \times 10^{-10} (L_{\text{IR}} + 2.2 L_{\text{UV}}) L_{\odot}$ . Bell et al. (2005) presents a near-identical calibration of  $\text{SFR}_{\text{IR+UV}}$ , which would yield identical values of  $\text{SFR}_{\text{IR+UV}}$ , within errors. The upper limits in the IR+UV star-formation rates indicate

sources for which the MIPS  $24\ \mu\text{m}$  flux has  $S/N < 3$ , and the  $3\sigma$  upper limit is adopted to estimate  $L_{\text{IR}}$ .

As shown in Figure 9, for four out of the five galaxies in our sample, the SFRs estimated from the SED and IR+UV are consistent with each other. Given that this includes three upper limits, however, this statement demonstrates only that the SED SFR estimates are not significantly higher than expected. COS-207144 is the only galaxy in our sample with an SFR derived from UV+IR much larger than the SFR from SED modeling. This result may be due to a difference in the UV and IR properties of the two galaxies that make up the source.

## 6. Discussion and Conclusions

In this work, we have presented NIR spectra of five ultra-massive galaxies in the redshift range of  $1.7 < z < 2.7$ . We determined a spectroscopic redshift for one of these galaxies

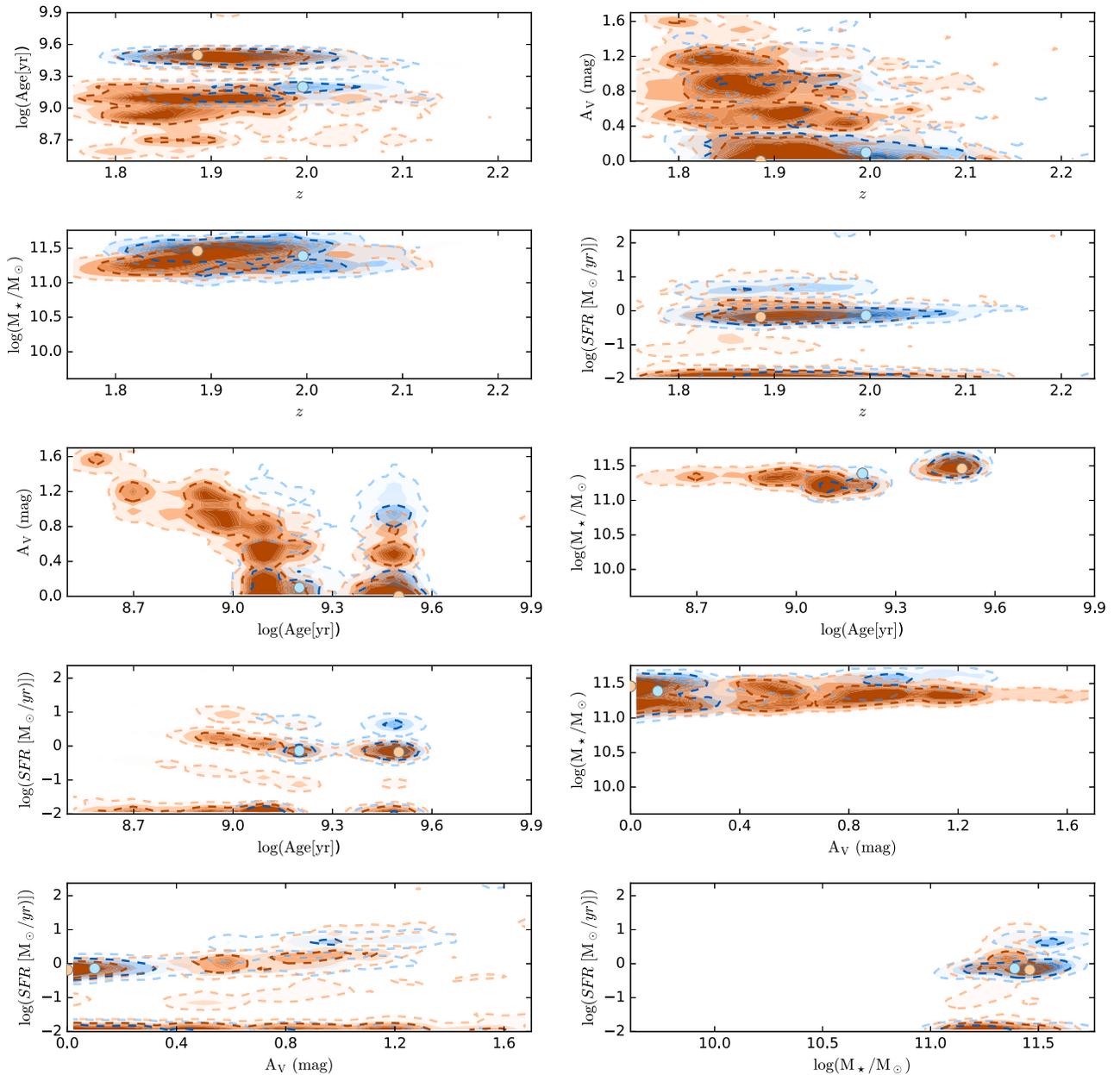


**Figure 12.** Same as Figure 10 for COS-37207.

based upon the presence of the Ca II H & K absorption features in its spectrum, and estimated continuum redshifts for the remaining galaxies in our sample using information from both broad- and medium-band UV-to-8  $\mu\text{m}$  photometry and the aforementioned spectra. The stellar population parameters of the galaxies were also estimated, confirming the ultra-massive status of these galaxies.

Out of the five galaxies, only one—COS-37207—is found to be star-forming, with  $\text{SFR} \approx 340 M_{\odot} \text{yr}^{-1}$ . This galaxy is also found to have the highest level of dust attenuation, with  $A_V \approx 1.70$  mag. A second galaxy, COS-71929, possesses properties that are intermediate with respect to the quiescent and star-forming populations at this redshift. The galaxy shows residual star-formation activity, i.e.,  $\text{SFR} \approx 9 M_{\odot} \text{yr}^{-1}$  and elevated dust obscuration, i.e.,  $A_V \approx 1.3$  mag, and sits well below the main sequence of star-forming galaxies at the redshift range of interest. The remaining galaxies in the sample

are classified as quiescent with substantially less dust extinction by both SED modeling and *UVJ* diagram classification. This finding is in agreement with the refined evolutionary path for the formation of local ultra-massive galaxies recently presented by Marchesini et al. (2014). Although the progenitors of today’s ultra-massive galaxies (i.e.,  $\log(M_{\star}/M_{\odot})$  11.8 at  $z \sim 0$ ) were massive, highly obscured, dusty star-forming galaxies at  $2.5 < z < 3$ , we expect the sample in hand to quench at earlier times on average, as the galaxies in the present sample have significantly higher stellar masses than the majority of the galaxies from Marchesini et al. (2014). Though the present sample is too small to estimate population characteristics, if we consider COS-37207 to be the only star-forming galaxy (grouping COS-71929, given its suppressed star formation, with the quiescent galaxies in the sample), we obtain a quiescent fraction of 0.80 for such galaxies ( $11.45 < \log(M_{\star}/M_{\odot}) < 11.9$ ) at  $1.7 < z < 2.7$ .



**Figure 13.** Same as Figure 10 for COS-207144.

There is no significant evidence that this quiescent fraction is systematically larger than the fraction of  $\sim 0.40$  estimated by Marchesini et al. (2014) for the population of the progenitors of today's ultra-massive galaxies at  $1.5 < z < 2.5$ . Marchesini et al. (2014) found a quiescent fraction of  $\approx 0.40$  for their sample at  $1.5 < z < 2.5$ , which covered a range of stellar masses of  $11.2 \lesssim \log(M_*/M_\odot) \lesssim 11.5$ . Similarly, one out of the three galaxies with stellar masses  $\log(M_*/M_\odot) \gtrsim 11.4$  in Belli et al. (2014) was found to be star-forming at  $2. < z < 2.5$ , and Muzzin et al. (2013b) estimates that the quiescent fraction should meet or exceed 0.50 at  $\log(M_*/M_\odot) \approx 11.3$  for the redshift range considered in this work.

Assuming a true quiescent fraction of 0.40, there is a significant ( $\approx 8\%$ ) chance that a sample of five galaxies will contain at least four quiescent galaxies if the sample was drawn without bias for the same distribution. However, given previous evidence that more massive galaxies enter quiescence at earlier

times relative to lower mass galaxies (see, e.g., Brammer et al. 2011; Muzzin et al. 2013b), we find that our data are consistent with these claims. However, in order to probe the true mass dependence of the quiescent fraction at a given redshift of such ultra-massive galaxies, a larger sample is needed in order to properly marginalize over confounding variables and to suppress the uncertainty due to low number statistics.

In order to better characterize the population of ultra-massive galaxies in the early universe, we must increase the number of identified ultra-massive galaxies and construct a larger number of well-sampled SEDs from which to derive stellar population property estimates and photometric redshifts. The fact that the inclusion of our NIR spectra, in the absence of a spectroscopic redshift, does not greatly affect the stellar masses and redshifts of the galaxies in our sample implies that such galaxies may be reliably identified using photometry alone, in agreement with Muzzin et al. (2009). The

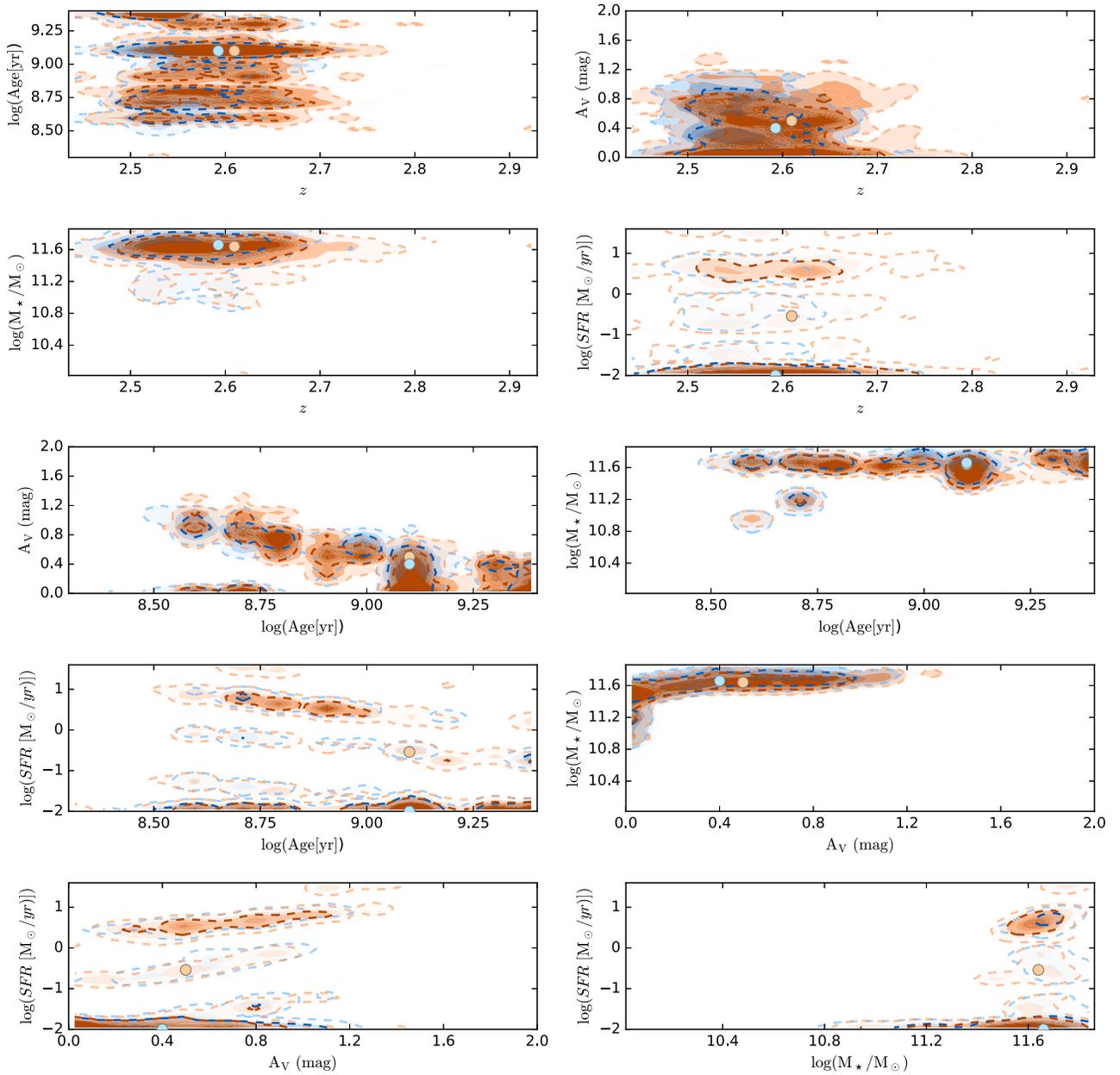


Figure 14. Same as Figure 10 for COS-90676.

NEWFIRM Medium-Band Survey II (NMBS-II) will identify these ultra-massive galaxies across  $5.2 \text{ deg}^2$  of the sky (the UltraVISTA DR1 catalog, upon which this work is based, spans  $1.62$  square degrees), and will provide medium-band photometry to produce well-sampled SEDs. Moreover, the VISTA VIDEO survey (Jarvis et al. 2013) includes oft-visited areas of the sky and covers nearly  $12 \text{ deg}^2$ , giving it both the wavelength range and spatial coverage to robustly identify many new ultra-massive galaxies.

Deeper spectroscopy will also greatly aid the study of this sample—the galaxies in hand are bright enough that deep spectroscopy is feasible from the ground, given an 8–10 m class telescope. Though our detection of absorption features in COS-71929 did not have a high-enough S/N to perform a detailed measurement of the velocity dispersion, preliminary estimates indicate a substantial velocity dispersion. Such

measurements would also be greatly helpful in constraining the systematics of SED modeling by providing independent measurements of output parameters.

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Based on data products from observations made with ESO Telescopes at the La Silla Paranal Observatory under ESO programme ID 179.A-2005 produced by TERAPIX and the Cambridge Astronomy Survey Unit on behalf of the Ultra-VISTA consortium. Based on observations obtained at the Gemini Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership: the National Science Foundation (United States), the National Research Council (Canada), CONICYT (Chile), Ministerio de Ciencia, Tecnología e Innovación Productiva (Argentina), and Ministério da Ciência, Tecnologia e Inovação (Brazil). This paper includes data gathered with the 6.5 m Magellan Telescopes located at Las Campanas Observatory, Chile.

This work made use of the products from the 3D-HST Treasury Program (GO 12177 and 12328) with the NASA/ESA *HST*, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS5-26555.

### Appendix Monte Carlo Parameter Correlations

Here we present the results of the Monte Carlo simulations to constrain the error on our stellar population property estimates (Figures 10–14). For each run, fluxes for both the spectroscopy and photometry are drawn from a normal distribution characterized by the error on the original measurement. Each instance is run with a fixed redshift drawn according to the best-fit EAZY redshift PDF, as shown in Figure 3.

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